

## The Lives of Stars

(pp. 462-473)

When you look at the stars on a clear night, the stars differ in \_\_\_\_\_. A \_\_\_\_\_ is a light collector, concentrating the light it collects in the \_\_\_\_\_. A \_\_\_\_\_ can be placed over the \_\_\_\_\_ and a long time exposure photograph can be taken. Astronomers have inferred a great deal from \_\_\_\_\_ about important properties of stars: \_\_\_\_\_, \_\_\_\_\_, \_\_\_\_\_, \_\_\_\_\_, and \_\_\_\_\_.

Although the human eye can only detect \_\_\_\_\_ light, other types of electromagnetic radiation also exist, each with a different range of \_\_\_\_\_. The entire arrangement of electromagnetic radiation is called the \_\_\_\_\_.

\_\_\_\_\_ and \_\_\_\_\_ can affect how bright a light-emitting object appears. This is also the case with stars. The closest stars to Earth are \_\_\_\_\_ the brightest because it depends both on how \_\_\_\_\_ the stars are to Earth and how \_\_\_\_\_ the stars are. \_\_\_\_\_ is a property of the stars themselves. Luminosity is a measure of the total amount of \_\_\_\_\_ a star radiates per \_\_\_\_\_.

Although the stars look like small points of \_\_\_\_\_ light, they range widely in \_\_\_\_\_. Some are \_\_\_\_\_ white, \_\_\_\_\_ green, yellow, or \_\_\_\_\_-red. Astronomers use the colour of a star to infer its \_\_\_\_\_. A \_\_\_\_\_ star, such as the Sun, is relatively \_\_\_\_\_, with a surface temp. of about \_\_\_\_\_. A \_\_\_\_\_ star is relatively cool, about \_\_\_\_\_. A \_\_\_\_\_ star is extremely \_\_\_\_\_, typically between \_\_\_\_\_ and \_\_\_\_\_. Analysis of starlight also indicates a star's \_\_\_\_\_. To analyze the light, astronomers use a \_\_\_\_\_ which separates light into a \_\_\_\_\_ (bands of different \_\_\_\_\_). A star's spectrum also shows dark \_\_\_\_\_ across the \_\_\_\_\_. These bands reveal that some \_\_\_\_\_ of light have \_\_\_\_\_, absorbed by \_\_\_\_\_ in the star's atmosphere. Each \_\_\_\_\_ (eg. \_\_\_\_\_, helium, \_\_\_\_\_, calcium, etc.) leaves its unique pattern of \_\_\_\_\_ on a spectrum. Astronomers use this information (and the basic temp.) to \_\_\_\_\_ stars.

Astronomers found they could calculate the \_\_\_\_\_ or \_\_\_\_\_ of a star once they knew its \_\_\_\_\_ and \_\_\_\_\_. Stars come in many sizes. More than half the stars we see from Earth are \_\_\_\_\_ stars, \_\_\_\_\_ single stars \_\_\_\_\_ one another. Star mass is expressed in terms of \_\_\_\_\_. The Sun is a \_\_\_\_\_ star; other stars range from \_\_\_\_\_ solar mass to over \_\_\_\_\_ solar masses.

The Hertzsprung-Russell diagram (Fig. 14.7) is a plot/graph of the \_\_\_\_\_ (on the y-axis) against the \_\_\_\_\_ or temperature (on the x-axis) of stars. The Sun is shown near the \_\_\_\_\_. A star located in the upper-right corner is \_\_\_\_\_ and \_\_\_\_\_ whereas a star in the lower-left corner is \_\_\_\_\_ and \_\_\_\_\_. About \_\_\_\_\_% of stars is not in the main sequence. These stars (such as \_\_\_\_\_ giants and supergiants) are very large but rather \_\_\_\_\_ in temperature. Other stars are very small but extremely \_\_\_\_\_ in temperature; these stars are called white \_\_\_\_\_.

\_\_\_\_\_ law of gravitation states that an \_\_\_\_\_ gravitational force exists between all \_\_\_\_\_ and this force gets \_\_\_\_\_ the closer the objects become. Gravity is the force that helps \_\_\_\_\_ and \_\_\_\_\_ stars and ultimately causes them to \_\_\_\_\_ (see Fig. 14.9). Astronomers speculate that vast clouds of gas and \_\_\_\_\_ called \_\_\_\_\_ are the birth places of stars. In these clouds, \_\_\_\_\_ pulls the "\_\_\_\_\_ - \_\_\_\_\_" material together. The accumulating gas causes the \_\_\_\_\_ in the centre to \_\_\_\_\_. When it reaches about \_\_\_\_\_, the nuclear reaction, \_\_\_\_\_ (transformation of \_\_\_\_\_ into \_\_\_\_\_)

begins and the star "turns on" or shines. Once the \_\_\_\_\_ process begins, it starts to consume the \_\_\_\_\_ fuel. \_\_\_\_\_ begins to accumulate in the \_\_\_\_\_ of the star. The interior of the star continues to heat up, increasing the \_\_\_\_\_ and \_\_\_\_\_. These forces are balanced by the \_\_\_\_\_ pull toward the centre. The result is a \_\_\_\_\_ star. All \_\_\_\_\_ sequence stars, including our \_\_\_\_\_, are in this condition. The time the star remains in this phase, before advancing to a later phase, depends on its \_\_\_\_\_.

Low mass stars ( \_\_\_\_\_ ) consume their hydrogen \_\_\_\_\_ over a period that may last as long as \_\_\_\_\_ years. During that time, they lose significant \_\_\_\_\_, essentially \_\_\_\_\_. In the end, all that remains of them is a very faint \_\_\_\_\_.

Intermediate mass stars (ie. the \_\_\_\_\_) consume their \_\_\_\_\_ a little \_\_\_\_\_, over a period of about \_\_\_\_\_ years. When the hydrogen is used up, \_\_\_\_\_ production stops and gravity causes the star to \_\_\_\_\_. As the core contracts, the \_\_\_\_\_ increases and the outer layers begin to \_\_\_\_\_. Eventually, the \_\_\_\_\_ starts fusing into \_\_\_\_\_. Because the star has expanded so much, the outer layers are much \_\_\_\_\_ than when the star was a main sequence star. It therefore appears \_\_\_\_\_, earning the name \_\_\_\_\_. Our Sun will evolve to this phase in about \_\_\_\_\_ years.

\_\_\_\_\_ winds peel away gases, eventually revealing the hot inner region of a star. The result is a \_\_\_\_\_, a fuzzy \_\_\_\_\_ object. Over time, it disperses into space, its remnant \_\_\_\_\_ slowly and losing \_\_\_\_\_. It then becomes a \_\_\_\_\_. These continue to cool over billions of years and are eventually nothing more than a dark cylinder called a \_\_\_\_\_.

High mass stars consume their \_\_\_\_\_ very rapidly. So much \_\_\_\_\_ is released that the star swells into a \_\_\_\_\_. In the core, \_\_\_\_\_ is transformed into \_\_\_\_\_. Once an \_\_\_\_\_ core is achieved, no further \_\_\_\_\_ process can counter the force of \_\_\_\_\_ and the core collapses. A massive \_\_\_\_\_ wave bursts from the star's surface causing a huge explosion or \_\_\_\_\_.

After the supernova phase, a star has one of two fates depending on its \_\_\_\_\_. If the remaining core is about 1.4 to 3 \_\_\_\_\_ masses, the gravity crushes the remnant into a small superdense object called a \_\_\_\_\_ star. If the core is greater than 3 \_\_\_\_\_ masses, they form \_\_\_\_\_ holes, objects so compact and \_\_\_\_\_ that not even \_\_\_\_\_ can escape. \_\_\_\_\_ holes are the most extreme conclusion of \_\_\_\_\_ work in stellar evolution.